

Synthesising, using, and correcting for telluric features in high-resolution astronomical spectra

A near-infrared case study using CRIRES*

A. Seifahrt^{1,2,3}, H. U. Käuffl³, G. Zängl^{4,5}, J. L. Bean², M. J. Richter¹, and R. Siebenmorgen³

¹ Physics Department, University of California, Davis, CA 95616, USA
e-mail: seifahrt@physics.ucdavis.edu

² Universität Göttingen, Institut für Astrophysik, Friedrich-Hund-Platz 1, 37077 Göttingen, Germany

³ European Southern Observatory (ESO), Karl-Schwarzschild-Str. 2, 85748 Garching, Germany

⁴ Ludwig-Maximilians-Universität München, Meteorologisches Institut, Theresienstrasse 37, 80333 Munich, Germany

⁵ Deutscher Wetterdienst, Frankfurter Str. 135, 63067 Offenbach, Germany

Received 1 December 2009/ Accepted 18 August 2010

ABSTRACT

We present a technique to synthesise telluric absorption and emission features both for in-situ wavelength calibration and for their removal from astronomical spectra. While the presented technique is applicable for a wide variety of optical and infrared spectra, we concentrate in this paper on selected high-resolution near-infrared spectra obtained with the CRIRES spectrograph to demonstrate its performance and limitation. We find that synthetic spectra reproduce telluric absorption features to about 2%, even close to saturated line cores. Thus, synthetic telluric spectra could be used to replace the observation of telluric standard stars, saving valuable observing time. This technique also provides a precise in-situ wavelength calibration, especially useful for high-resolution near-infrared spectra in the absence of other calibration sources.

Key words. atmospheric effects – instrumentation: adaptive optics – instrumentation: spectrographs – methods: observational – methods: data analysis

1. Introduction

Ground based astronomical observations suffer from the disturbing effects of the Earth's atmosphere. Strong ro-vibrational line systems of water vapour (H₂O), carbon dioxide (CO₂), and ozone (O₃) shape the observational windows in the near infrared that define the bandpasses of the *J*, *H*, *K*, *L*, and *M* filters from 900 to 5500 nm. In addition, nitrous oxide (N₂O) and methane (CH₄) contribute noticeably to the atmospheric transmission losses in the near infrared, hampering the observation of important astrophysical lines.

The Earth's atmosphere below about 60 km is in local thermodynamic equilibrium and Kirchhoff's law applies. While the transitions of the molecules discussed above appear in absorption in a stellar spectrum, they also appear in emission on the sky since the Earth's atmosphere is radiating its thermal energy into space in these lines. The emission originating from the Earth's atmosphere competes with the astronomical signal at wavelengths longer than about 2300 nm, depending on ambient temperature, slit width, and the readout noise of the detector. Because their spatial signature is extended and moderately flat, beam-switching (nodding) techniques or in-situ fitting and subtraction in 2D spectra is generally sufficient to remove the direct influence of the emission lines in the near and mid-infrared.

The removal of the absorption features is considerably more difficult because no in-situ reference for these features is available in the optical and near-infrared. The usual technique is to observe a star with either very few or well determined intrinsic

features (a “telluric standard”) close in time and airmass to the scientific target. The choices for such a standard are early-type stars (spectral types mid B to late A) that exhibit a rather clean continuum despite strong hydrogen lines, or solar-like stars of spectral type early to mid G. In the latter case, high resolution Fourier transform spectra (e.g., Hase et al. 2010; Livingston & Wallace 1991) of the sun provide the necessary template to correct for the intrinsic stellar lines of the standard star.

Despite the wide use of this empirical calibration (e.g., Vacca et al. 2003, and references therein), the use of standard stars for removing telluric absorption lines has several disadvantages, especially when the correction is needed for low resolution spectra. At spectral resolving powers exceeding $R \sim 20\,000$, telluric lines start becoming resolved and bandpasses with high telluric line density become accessible to ground based observations. The full potential of high resolution spectroscopy can only be used when unsaturated telluric lines are removed without compromising the intrinsic signal-to-noise ratio (*SNR*) of the data. To reach noise limited performance, the path length (airmass) and instrumental profile between the standard star and the science target must match. Also, the standard must be bright enough in order not to limit the final *SNR*.

The intrinsic lines of early type standard stars, even when being accounted for, make the precise study of near-infrared hydrogen lines very difficult (Vacca et al. 2003) and the choice of a fast rotator can be an advantage. Also, the mismatches in the line depths of solar-type standard stars with the solar FTS atlas, caused, e.g., by differences in metallicity or effective temperature, hamper precise abundance studies from near-infrared metal lines. Another problem arises for the *L* and *M* band, where the source flux of early- and solar-type standard stars can be the

* Based on observations made with ESO Telescopes at the Paranal Observatory under programme IDs 60.A-9076(A), 079.D-0810(A), and 179.C-0151(A).

limiting factor, favoring the observation of non-stellar sources, like bright asteroids.

Additional limitations in the use of empirical telluric corrections may be based on the spectrograph itself. Limited spectral coverage around the hydrogen lines of early-type standard stars prohibit fitting and correcting these lines, which leaves a disturbed continuum in the spectrum of the science target. Adaptive optics fed long-slit spectrographs, such as CRIRES at the VLT (see Sect. 2), introduce another level of complication. The diffraction-limited full width at half maximum (*FWHM*) of the stellar point spread function (PSF) can be considerably smaller than the slit width, so that one often observes with an under-filled entrance slit. The actual instrumental profile, and hence spectral resolving power¹, is then governed by the PSF of the source, rather than by the slit width of the spectrograph. Moreover, slight displacements of the source in the slit will lead to shifts in the wavelength at the detector plane. In the case of CRIRES, a shift of only 30 m/s (1/50 of a CRIRES pixel or 4 mas at the entrance slit) can notably degrade the telluric correction (see, e.g., Sect. 4.1.2). Since most of the telluric lines in the near-infrared are not fully resolved, even at $R = 100\,000$, their shapes are governed by the instrumental profile. As a consequence, even at the same airmass the telluric standard star cannot perfectly correct the telluric features in the spectrum of the science target because differences in the performance of the adaptive optics ultimately lead to differences in the instrumental profile and variations of the wavelength solution. A further drawback in the use of telluric standard stars is the waste of observing time for the acquisition and observation of these stars, especially under the constraint of matching airmasses at small time differences. As mentioned earlier the problem of source brightness becomes especially apparent at longer wavelengths where one can easily spend as much or even more time observing a telluric standard than a science target.

While telluric emission and absorption features have been discussed so far as a disturbance in astronomical spectroscopy, they can as well turn out to be very useful. Determining the wavelength solution of near infrared spectra, especially at high spatial and spectral resolution, proves to be very difficult. Rare gas emission line lamps used in low resolution near-infrared spectrographs, such as He, Ne, Xe, and Kr lamps, provide much too sparse of a line coverage to be useful at high spectral resolution. Even the rich spectra of ThAr emission line lamps, commonly employed by high resolution optical spectrographs, have a much lower line density in the near-infrared than in the red optical. Typical line densities of ~ 400 lines per $\Delta\lambda = 100$ nm around $\lambda = 1000$ nm drop quickly to less than ~ 20 lines per $\Delta\lambda = 100$ nm at $\lambda = 2500$ nm (Kerber et al. 2008). Given the typical wavelength coverage of $\sim \lambda/200$ per detector, many CRIRES spectral settings remain poorly calibrated. Moreover the dynamic range of the ThAr emission spectrum is very high, owing to the strong contrast between the Th and Ar lines, leaving weaker lines at low SNR while nearby stronger lines saturate quickly.

Recently, frequency combs (Udem et al. 2002) have been proposed for the ideal spectral calibration of high resolution spectrographs. In principle a frequency comb would provide for high line density while linking the dispersion solution of the spectrograph to the time standard. However, the line density of a simple comb is much too high, so that some optical filtering is required, which will produce detrimental artifacts. Also, a

¹ Throughout this paper we define the spectral resolving power as $R = \lambda/\Delta\lambda$, where $\Delta\lambda$ is the *FWHM* of the instrumental profile.

frequency comb might solve for the problem of insufficient line density and uniformity of emission line lamps but the instrumental profile will still be different than for the science target, ultimately limiting the final precision of the wavelength calibration of the science spectrum.

This limitation applies also for atmospheric emission lines from ro-vibrational levels of the Hydroxyl radical OH* (Meinel bands, commonly referred to as OH* airglow emission) which are routinely used for the calibration of low-resolution near-infrared spectrographs (e.g., Rousselot et al. 2000). Moreover, neither the line strength nor the line density is sufficient to be used as a useful wavelength calibration over a wide spectral range at high spectral and spatial resolution, as provided by CRIRES. The small spatial pixel scale of CRIRES (86 mas/pix) makes the use of OH* lines for calibration purposes challenging, given that each pixel is less than 10^{-2} arcsec² on the sky. In contrast, the atmospheric absorption and emission features of other molecules show a high density of strong absorption lines and thus provide a natural in-situ wavelength calibration. Moreover, the telluric absorption lines suffer the same instrumental effects as the lines in the stellar spectra and, hence, provide an intrinsically more precise calibration source as compared to telluric emission lines or lamp emission lines, the latter even recorded in separate exposures and often hours apart from the science spectra.

We have thus investigated the feasibility of calculating theoretical atmospheric transmission spectra in order to use them for wavelength calibration and to remove their signature from the astronomical spectra, thus, replacing standard stars as a source of telluric calibrators. This idea is not new and has been used successfully before (Lallement et al. 1993; Widemann et al. 1994; Bailey et al. 2007). Nevertheless, modeling telluric spectra has not evolved to a standard technique in optical and infrared spectral data reduction because suitable radiative transfer codes were hard to access and synthesising adequate spectra was hindered by the incompleteness of molecular line databases. This situation has improved in the last years.

In an effort to characterise the abilities and the efficient usage of the CRIRES spectrograph, we have developed a general method to model the transmission and emission spectrum of the Earth's atmosphere above Cerro Paranal. This method can be used to perform a wavelength calibration of CRIRES and to subsequently remove the telluric absorption lines apparent in these spectra. Moreover the model can also be used to predict the performance of planned CRIRES observations using a forecast of the atmospheric conditions during a planned observing run.

The paper is structured as follows: in Sect. 2 we describe the data reduction of the CRIRES observations. In Sect. 3 we describe the radiative transfer code and input parameters used to synthesise the telluric spectra, which we compare to the observations in Sect. 4. We summarise our findings in Sect. 5 and close with conclusions about the performance of the presented technique.

2. Observations and data reduction

CRIRES is a high resolution ($R \leq 100\,000$) near-infrared ($\lambda\lambda = 960\text{--}5200$ nm) adaptive optics (AO) – fed spectrograph at the VLT on Paranal, Chile (Käufl et al. 2004, 2006a,b; Paufigue et al. 2006). Commissioning and science verification (SV) observations were executed in October 2006 and February 2007. The spectrograph is available to the community and has been used for regular observations since April 2007.

Table 1. Log of CRIRES observations retrieved from the ESO archive.

Object	Spectral-type ^a	Date and time	λ_{ref} (nm)	ν (GHz)	Slitwidth ($''$)	R^b	Air-mass	SNR^c
HD 74575	B1.5III	2007-05-10 01:00 UT	1134.0	264 367	0.2	64 000	1.30	250
HD 148605	B3V	2007-03-03 09:30 UT	1504.8	199 224	0.2	80 000	1.03	200
HD 71155	A0V	2007-02-28 04:30 UT	2076.5	144 374	0.4	60 000	1.24	300
HD 121847	B8V	2008-02-29 08:30 UT	2336.2	128 325	0.4	65 000	1.01	300
HD 118716	B1III	2007-03-05 05:00 UT	3907.8	76 716	0.2	65 000	1.30	380
HD 177756	B9Vn	2007-04-22 10:00 UT	4760.8	62 971	0.2	87 000	1.06	200

Notes. ^(a) SIMBAD spectral type. ^(b) Estimated resolving power from fitting telluric lines in the analysed spectral region. ^(c) Estimated signal-to-noise ratio in the continuum of the analysed spectral region.

We selected observations from both SV periods as well as more recent observations from the ESO data archive to test the telluric model spectra described in Sect. 3. For this purpose we chose observations of standard stars at different wavelengths, observed at different airmasses with nominal spectral resolving powers of $R = 50\,000$ and $100\,000$. The log of observations is given in Table 1.

Observations were always obtained in an AB or ABBA nodding pattern. Data reduction followed the standard steps for long-slit spectrographs as applicable to the CRIRES data format. All raw frames were treated with a non-linearity correction before pairwise subtraction removed the atmospheric emission features. The individual A-B and B-A frames were then divided by a normalised flatfield. Optimally extracted 1D spectra at both nodding positions were obtained by using a custom made IDL script based on Horne (1986).

Due to the curvature of the slit, it is inadvisable to combine the 2D spectra before the extraction and wavelength calibration step as is done in the current version of the CRIRES data reduction pipeline. The amount of curvature can be of the order of 1 pixel between two nodding positions at a nod throw of $10''$ and is variable in both spatial and spectral directions. We thus combined the individual 1D spectra after correcting for the slit curvature by producing a telluric model spectrum using the method outlined in Sect. 3 and interpolating the wavelength calibrated spectra to a common wavelength vector. All reduction steps were performed individually for the four chips of the CRIRES focal plane array.

3. Spectral synthesis

To compute theoretical transmission and radiance spectra of the Earth's atmosphere, a radiative transfer code is used that takes as an input a model of the atmosphere in terms of vertical temperature, pressure and molecular abundance profiles, as well as line data (frequency, line strength, pressure broadening coefficients, etc.). In this section we describe the usage of such a code as well as the necessary inputs in further detail.

3.1. Radiative transfer code

There are several radiative transfer codes tailored for the construction of telluric spectra. These codes are based on line-by-line computations of a layered model of the Earth's atmosphere.

STRANSAC (Scott 1974) is one the oldest of such codes. It utilises the GEISA spectral line catalogue (Jacquinet-Husson et al. 1999, 2003).

The 4A (Automatized Atmospheric Absorption Atlas) code also uses GEISA, and it provides a fast and accurate

line-by-line radiative transfer model that is particularly efficient in the infrared region between 3000 nm and 17 000 nm. The latest version of 4A can be retrieved free of charge for non-commercial usage from NOVELTIS², France.

ATRAN (Lord 1992) is another code to compute synthetic spectra of atmospheric transmission. It is based on a fixed atmospheric layering (see Sect. 3.2) and uses the HITRAN database (see Sect. 3.3) for line data input.

RFM (Reference Forward Model) is a line-by-line radiative transfer model originally developed at Oxford University³, under an ESA contract to provide reference spectral calculations for the MIPAS instrument launched on the ENVISAT satellite in 2002. The code uses HITRAN 2000 as its line database.

In this paper we only consider FASCODE (Clough et al. 1981, 1992), and used both its commercial version FASCODE3P/PcLnWin⁴ and the free of charge version LBLRTM⁵. LBLRTM is available as FORTRAN source code and can be compiled on various platforms. Its most common application is the spectral retrieval from ground and satellite based measurements in climatology. The code accepts a custom model atmosphere and uses the HITRAN database (see Sect. 3.3) as input for the line data. We found LBLRTM easy-to-use, yet flexible, and computed all spectra presented in this paper with this code.

3.2. Input I: model atmosphere

The choice of a representative model atmosphere is crucial for a realistic theoretical spectrum of the atmosphere. One of the more basic descriptions of a model atmosphere is given by the 1976 US Standard Atmosphere⁶, and its geographical subdivisions (tropical, mid-latitude and sub-arctic summer and winter models). Due to the geographical location of Cerro Paranal and the fractional changes in the abundance of trace gases since its definition, the 1976 US Standard Atmosphere gives a rather poor representation of the atmospheric conditions above Cerro Paranal. For example, the average surface temperature of the model at 2640 m is only 271 K and thus too low by about 10 K on average.

In addition to the semi-static distribution of atmospheric trace gases, one has to account for the slowly variable temperature and pressure levels as well as the highly variable water vapour content. Thus, one has to base the calculations on a more

² <http://www.noveltis.net/4AOP/>

³ <http://www.atm.ox.ac.uk/RFM/>

⁴ <http://www.ontar.com/Software/Products.aspx>

⁵ http://rtweb.aer.com/lblrtm_description.html

⁶ http://ntrs.nasa.gov/archive/nasa/casi.ntrs.nasa.gov/19770009539_1977009539.pdf

Table 2. Precipitable water vapour values (in mm PWV at zenith).

Date and time (UT)	GDAS		MM5		ESO ^a
	predicted	fitted	predicted	fitted	predicted
2006-10-10 03:00	1.8	0.8	1.8	0.8	1.9
2007-03-01 06:00	5.3	2.6	–	–	3.1
2007-03-03 09:00	4.1	1.6	2.4	1.6	1.2
2007-03-05 06:00	10.8	8.6	–	–	3.9
2007-04-22 09:00	9.6	7.2	–	–	2.5
2007-05-10 01:00	2.0	0.4 ^b	–	–	3.0
2008-02-29 09:00	3.6	2.2	–	–	2.8

Notes. ^(a) Prediction for Paranal Observatory, based on satellite measurements. Obtained from <http://www.eso.org/gen-fac/pubs/astclim/forecast/meteo/ERASMUS/>. ^(b) PWV determined from water vapour lines not shown in Fig. 10.

flexible and realistic model atmosphere which takes meteorological inputs into account.

We consider here the use of meteorological models from the Air Resources Laboratory (ARL) at the US National Oceanic and Atmospheric Administration (NOAA). Sounding files for temperature, pressure and dew point temperature in the troposphere and lower stratosphere (surface height ≤ 26 km) are available at the ARL website⁷ based on either Global Forecast System (GFS) or archived Global Data Assimilation System (GDAS) models. GFS analysis data are available in 3 h intervals for the next 84 h and 12 h intervals for the next 384 h. GDAS models are available in 3 h intervals for all dates since Dec. 2004.

The horizontal resolution of the GFS and GDAS models is 1° (approx. 110 km). Given that Cerro Paranal is very close to the Pacific ocean (~ 12 km), the influence of the ocean climate on the chosen grid point could compromise its validity. We have thus also computed model atmospheres for three nights during CRILES SV periods on October 10, 2006, and March 3–4, 2007, based on the Fifth-Generation NCAR/Penn State Mesoscale Model (MM5) (Grell et al. 1995), which uses the operational analyses of the European Centre for Medium-Range Weather Forecasts (ECMWF) as start and boundary conditions and was refined for the local topography of the Paranal area with a final resolution of 1 km in the inner integration area. Both meteorological models reproduced the ambient pressure and temperature at the VLT as recorded in the file headers of the analysed observations to within 5 hPa and 1.5 K, respectively.

All atmospheric models have to be supplemented with information on the vertical distribution of all molecules other than H₂O and used in the radiative transfer calculations. Moreover, since the meteorological models are limited to estimated surface heights of ~ 26 km, we also have to add temperature, pressure and water vapour profiles for the remaining atmospheric layers ($26 \leq h \leq 75$ km). For this purpose we have chosen an equatorial MIPAS model atmosphere⁸, constructed by John Remedios (U. Leicester). This model contains temperature, pressure and abundance information for 30 molecular species in 121 levels from 0–120 km at 1 km spacing. In Fig. 1 we show the mixing ratios of the 10 most abundant molecules in this model. The actual amount of molecules per cm³ drops roughly exponentially with the surface height and the predominant part of the telluric transmission is shaped for most species in the first few km above the observatory.

In Fig. 2 we compare our modified MM5 model to a GDAS model for March 3, 2007 at 09:00 UT. The MM5 was used with a set of modifications developed by Zängl (2002, 2003) in order to improve the numerical accuracy over steep topography. While temperature and pressure profiles are nearly identical, the water

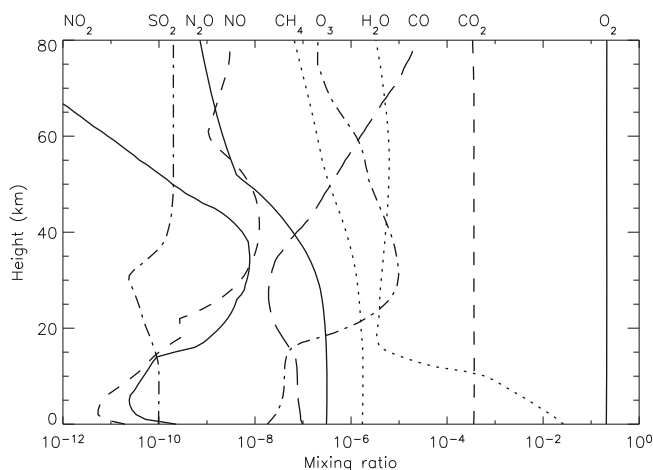


Fig. 1. Vertical profiles of the 10 most abundant molecules included in our model atmosphere. The water vapour (H₂O) abundance for the lowest atmospheric layers (surface height ≤ 26 km) is replaced by a meteorological model (GDAS or MM5) for the time of observation.

vapour profile is noticeably different. The GDAS model predicts a total water vapour column, PWV⁹, of 4.1 mm, the modified MM5 model predicts 2.4 mm. Moreover, the vertical distribution of the water vapour is different as well. The GDAS model predicts a much more humid surface layer, but dries out beyond the prediction of the MM5 model above 6 km height. A similar situation was found for October 10, 2006. Both models predict the same PWV of 1.8 mm but the GDAS model is again more humid below 3 km and less humid above. In any case, the models are often too wet when compared to the observations. The total water vapour amount was always scaled by 40–70% to match the observed line depth (see Sect. 4). The different vertical distributions of the water vapor did not lead to a significant difference in the fit quality of the water vapor lines once the total water vapor amount was rescaled to reach the best fit. We compare the predicted and fitted PWV values in Table 2 and conclude here, that the GDAS models are applicable and sufficient as meteorological input to refine the global equatorial model atmosphere. See Sect. 4 for a further discussion about the performance and limitation of the actual fitting.

3.3. Input II: molecular line database

HITRAN, in its latest 2008 edition (Rothman et al. 2009), is a molecular line database containing frequency, line strength and pressure broadening coefficients for more than 1.7 million spectral lines of 42 different molecules and their common isotopes. The database is regularly updated and supplemented with new

⁹ Measured as the precipitable water vapour amount at zenith.

⁷ <http://www.arl.noaa.gov/ready/cmet.html>

⁸ <http://www-atm.physics.ox.ac.uk/RFM/atm/>

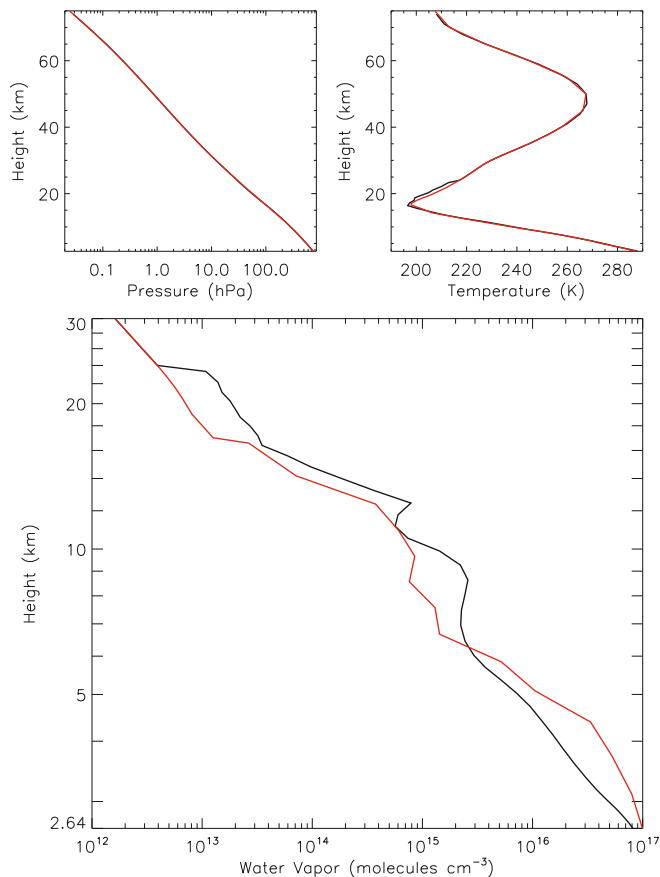


Fig. 2. Comparison of atmospheric models for Cerro Paranal, GDAS (red) and MM5 (black). Shown are pressure (*upper left panel*), temperature (*upper right panel*) and water vapour (*lower panel*) profiles for March 3, 2007 at 09:00 UT for surface heights above Paranal (2.64 km). Information for heights greater than 26 km (GDAS) and 22 km (MM5) were delivered by a global model for equatorial regions.

line data and molecular cross-sections. LBLRTM is distributed with a line file that is based on the HITRAN 2004 version and includes individual updates for some molecules until June 2009. A complete update to HITRAN 2008 is planned for the near future (Mark Shephard, private communication).

4. Results

We have computed synthetic telluric transmission spectra for a number of selected CRILES observations (see Table 1) using LBLRTM and a GDAS model atmosphere closest in time to the observation. The wavelength range of the model was set to the coverage of the chosen CRILES setting. LBLRTM uses the altitude angle of the telescope pointing, the height of the observatory and the model atmosphere to calculate the line-by-line and layer-by-layer transmission and radiance. We make no use of the calculated radiance at this point, since the telluric emission is removed conveniently by nodded observations, and use the transmission spectrum only.

An IDL based Marquardt-type χ^2 minimisation code (Bevington & Robinson 1992, p. 161) was used to (a) refine the wavelength solution by fitting a second order polynomial, using the original wavelength calibration as first guess; (b) re-determine the total water vapour amount and column density of other molecules present in the chosen setting; and (c) determine the instrumental profile and continuum rectification parameters.

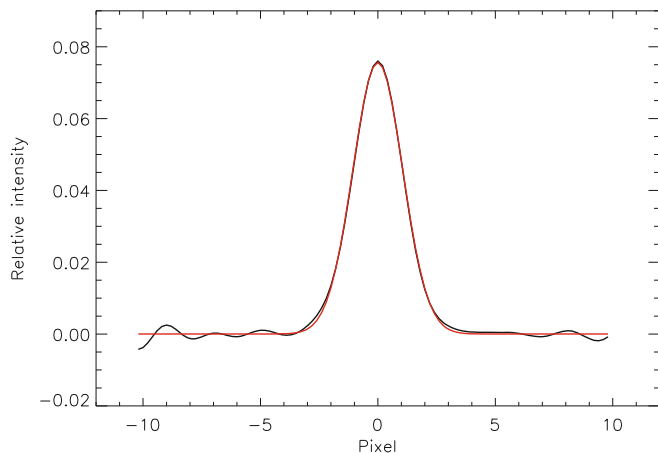


Fig. 3. CRILES instrumental profile, obtained from SVD fitting of methane (CH_4) lines in setting λ_{ref} 2336 nm (see Fig. 6). Black: empirical profile constructed from 21 Eigenvalues. Red: best fit Gaussian model with a *FWHM* of 2.67 pixel (4.0 km s^{-1} , $R = 70\,000$).

The determination of the instrumental profile (IP) is crucial for a proper fitting of the only partially resolved atmospheric lines. We have used a singular-value decomposition (SVD) technique in an implementation given by Rucinski (1999) to determine the instrumental profile. The wavelength scale was interpolated to a five times higher sampling than the original, Nyquist sampled wavelength scale and a 20 pixel wide instrumental profile was constructed from 11–13 Eigenvalues from all telluric lines present in the spectrum. Higher number of Eigenvalues did not lead to a further improvement of the fit but amplified the noise in the instrumental profile. In all cases, the instrumental profile was extremely well matched by a single Gaussian function whose *FWHM* depends on the slitwidth and the performance of the adaptive optics, see Fig. 3.

In cases where extremely narrow spectral lines were present (e.g. the unresolved ozone lines at 4700 nm – see Sect. 4.1.5) up to 30 Eigenvalues could be used without amplifying the noise. This analysis confirmed that a Gaussian model is a valid approximation for the instrumental profile at signal-to-noise levels below ~ 300 . We have thus used a single Gaussian model for the final fitting of the telluric model, the width of the profile being the only free parameter. We note that our determined *FWHM* are broader than expected for $R = 100\,000$. At the time most of the data presented in this paper were recorded, the focal plane array of CRILES was slightly out of focus. Thus, the maximum resolving power measured in the presented dataset is $R \sim 87\,000$.

4.1. Performance in selected settings

We present CRILES observations of bright standard stars that contain clear and isolated line systems of the most abundant molecules of the Earth’s atmosphere. We evaluate the performance of the telluric correction by checking the residuals for systematics by a rigid quantisation of the signal-to-noise ratio (SNR) in the corrected spectra. We chose the SNR as a measure for the performance of the telluric model fit since it has the strongest practical implication for the observer, as it allows a judgement as to which precision a position or an equivalent width of a stellar line can be measured in a spectrum. In regions of telluric contamination, the residuals from the removal of telluric lines are the main factor hindering a high signal-to-noise measurement, otherwise limited only by the source flux or the uniformity of the detector response.

In order to compare the SNR before and after the correction, we first compute the intrinsic noise level of our CRILES spectra by adding the photon noise of the source and the measured background noise in quadrature for each pixel along the dispersion direction. The background noise consists of the detector readout noise, the photon noise of the dark current and the photon noise of the thermal emission of the atmosphere and the telescope after pairwise subtraction of nodded frames. We have validated this procedure by measuring the standard deviation in the continuum level of each spectrum. The result was always consistent with our noise estimate.

Hence, for the input spectra we have the SNR as a function of wavelength which scales with the square root of the received flux, and thus with the square root of the atmospheric transmission T . In addition, the telluric emission lines in the thermal infrared example ($\lambda \sim 4700$ nm, Sect. 4.1.5) add noise where the lower transmission of the same telluric lines in absorption degrade the SNR in comparison to the continuum level.

The fitting procedure minimises the residuals in the observed-computed spectrum (O-C) over the whole wavelength region of the considered spectral range of $\sim \lambda/200$, recorded on one CRILES detector. For the removal of the telluric lines, however, the measured spectrum has to be divided by the computed model. After division by the model, we measure the noise in the spectrum as the standard deviation of all pixels. The SNR is then simply the inverse of the measured noise, given that after the division the signal level is at unity for a normalized spectrum.

It is important to remember that even for a perfect model of the telluric lines the O/C residuals will be larger, and the SNR smaller, in comparison to the undisturbed continuum, since the spectrograph receives less flux where telluric absorption lines are present. Hence, when dividing through the telluric transmission spectrum, and thus correcting the received flux for the presence of these lines, the noise in these regions is higher than in the neighboring continuum, simply because less flux was received in the lines in the first place. Albeit this caveat, we use the formal SNR as a direct measure of the performance of the modeling since it reflects the impact of the telluric lines on the astronomical spectrum and gives the observer a meaningful number to work with.

4.1.1. H₂O at 1505 nm

Water vapour (H₂O) is the most important absorber in the near-infrared. In addition to the strong ν_1 and ν_3 fundamental bands at 2700 nm, several overtone and combination bands centered at 940 nm, 1100 nm, 1380 nm, 1870 nm, and 3200 nm are present in near-infrared spectra. We picked a few isolated, unsaturated H₂O lines at 1505 nm in a CRILES standard star setting from March 03, 2007. The airmass of the observation was 1.033, the resolving power was $\sim 80\,000$ at a slitwidth of $0.2''$. Both the GDAS model and the modified MM5 model had to be scaled to a PWV of 1.6 mm for the best fit with the observation (see Fig. 4). The achieved correction using the GDAS and MM5 models is generally satisfying but falls behind the expectation from the noise limit. The average SNR of the measured spectrum is ~ 200 . After the division by the model spectrum, the average SNR drops to ~ 180 . In the line centers (40–70% transmission) the SNR drops from ~ 160 to ~ 50 after dividing by the model. Note how lines of equal strength, e.g., around 1504 nm are fitted with greatly varying precision, pointing to uncertainties in the line parameters.

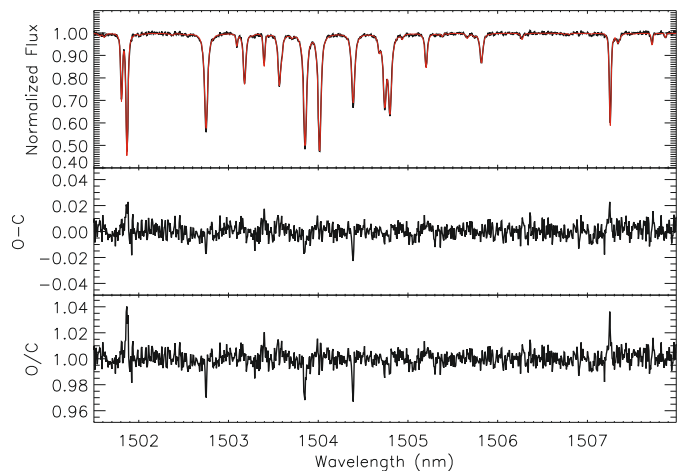


Fig. 4. Water vapour (H₂O) lines on chip three of CRILES setting λ_{ref} 1504.8 nm. *Top panel:* CRILES spectrum (black), and synthetic transmission spectrum (red). *Lower panels:* O-C and O/C residuals of the fit.

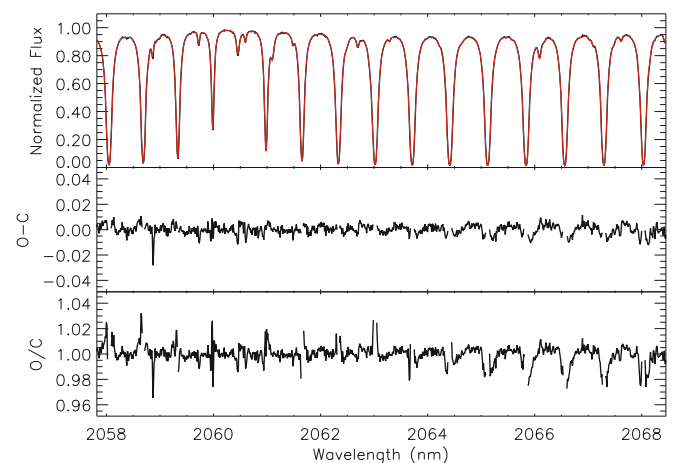


Fig. 5. Same as Fig. 4 but for carbon dioxide (CO₂) lines on chip two of CRILES setting λ_{ref} 2076.5 nm. Regions with telluric transmission of less than 5% are not shown in the residuals.

4.1.2. CO₂ at 2060 nm

Carbon dioxide (CO₂) has a nearly constant vertical mixing ratio (see Fig. 1) but due to its strong interaction with the biosphere, the low altitude CO₂ abundance can be expected to be variable in time. CO₂ is the third most abundant molecule after oxygen and nitrogen at heights exceeding ~ 4 km and exhibits various strong absorption bands centered at 1400 nm, 1600 nm, 2000 nm, 2700 nm, and 4300 nm. Kenworthy & Hanson (2004) explored the performance of empirical telluric line removal around the important He I lines at 2058.69 nm, which is buried under a comb of lines from the R branch of the 20013 \leftarrow 00001 transition of telluric $^{12}\text{C}^{16}\text{O}^{16}\text{O}$.

We have used a standard star observation obtained on March 01, 2007 in this spectral region to test the performance of our synthetic telluric model. The airmass of the observation was 1.24, the resolving power was $\sim 60\,000$ at a slitwidth of $0.4''$. The CO₂ abundance is given by the global atmospheric model (see Sect. 3.2) and had to be scaled down by $\sim 7\%$ to get the best match with the observation (see Fig. 5), despite the fact that the atmospheric CO₂ abundance is continuously growing and shows yearly changes of only $\pm 1\%$ (NOAA/ESRL¹⁰).

¹⁰ www.esrl.noaa.gov/gmd/ccgg/trends/

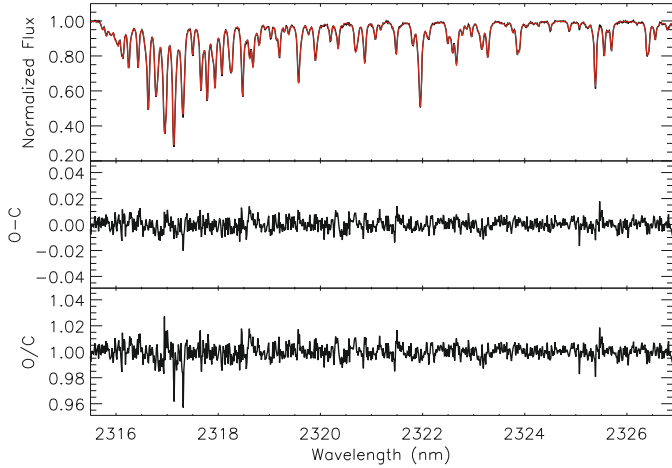


Fig. 6. Same as Fig. 4 but for methane (CH_4) and a few water vapour (H_2O) lines on chip two of CRIRES setting $\lambda_{\text{ref}} 2336.2$ nm.

The average SNR of the spectrum is ~ 300 . It drops to about 230 at transmissions $\geq 60\%$ when corrected for the telluric absorption. The SNR at the line cores (10–50% transmission) is ~ 160 and drops to ~ 50 after the correction, see Fig. 5. As already noted by Kenworthy & Hanson (2004), a precise wavelength solution is crucial for a good fitting. Wavelength errors at the level of 1/50 pixel (~ 30 m/s) notably degrade the fitting quality because of the steep line profiles. While the line positions in the HITRAN database have uncertainties at the m/s level, the uncertainty in the pressure shifts may become an issue at this level of precision.

After dividing by the best fit synthetic model, the region around the He I line has a SNR of ≥ 160 . We note that a high- J line of $^{13}\text{CO}_2$ leaves a 3% residual right at the rest wavelength of the He I line. We also note that for increasing J quantum numbers, the line shape of the $^{12}\text{CO}_2$ is less well matched as close to the band center. CO_2 exhibits strong line coupling which is taken into account by LBLRTM (see discussion in Sect. 5). It remains thus unclear whether the systematic increase of the residuals are due to uncertainties in HITRAN or due to insufficient treatment of the line coupling.

The $^{12}\text{CO}_2$ lines are fully resolved and only a few narrow lines are present in the setting. This inhibits a precise determination of the instrumental profile, which becomes very sensitive to noise in the spectrum. Thus, the width of the instrumental profile is partly degenerate with the line strength of CO_2 in the model. The exact locus of the best correction can be modified by slight changes of the $FWHM$ of the instrumental profile. This favours a customised telluric correction, tailored for a specific stellar line.

4.1.3. CH_4 at 2320 nm

Methane (CH_4) exhibits strong bands centered at 1660 nm, 2200 nm, and between 2300 nm to 3830 nm, although it is much less abundant than H_2O and CO_2 . The methane abundance in our model is set by the global atmospheric model (see Sect. 3.2). Methane is another of the so called *greenhouse gases*. Its atmospheric concentration is slowly increasing with time, but its growth rate dropped considerably in the last years (Dlugokencky et al. 2003). We chose an observation from Feb. 29, 2008, covering a strong CH_4 band at 2320 nm for fitting, see Fig. 6. The airmass of the observation was 1.01, the resolving power was $\sim 65\,000$ at a slitwidth of $0.4''$. This region is populated by CO first overtone lines in astrophysical objects and thus marks a

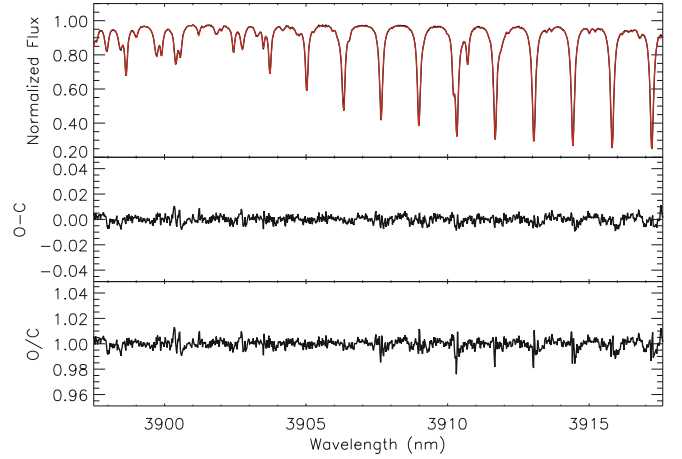


Fig. 7. Same as Fig. 4 but for nitrous oxide (N_2O) lines on chip three of CRIRES setting $\lambda_{\text{ref}} 3907.2$ nm. Minor contributions between $\lambda\lambda 3898$ – 3905 nm arise from H_2O and CH_4 .

very interesting spectral region. Besides the strong CH_4 lines, there are also three prominent water vapour lines in the covered region, one strong line at 2321.95 nm, and two weaker line blends at 2322.66 nm and 2323.27 nm. The best fit PWV at zenith was 2.2 mm. We did not have to scale the original CH_4 abundance in the equatorial model. The average SNR of the spectrum is ~ 300 , dropping to ~ 220 when corrected for the telluric absorption.

4.1.4. N_2O at 3907 nm

An observation from March 05, 2007 covered a nitrous oxide (N_2O) band around 3907 nm at airmass 1.3. The resolving power was only $\sim 65\,000$ at a slitwidth of $0.2''$. N_2O is a factor of two less abundant than CH_4 . Still, it exhibits strong line systems between 2870 and 4500 nm. As for CO_2 , we had to scale the N_2O abundance in our model down by $\sim 5\%$ to match observations. In addition to the strong N_2O lines, there are weaker lines of CH_4 and H_2O present in this setting, see Fig. 7. The best fit PWV at zenith was 8.6 mm. The average SNR of the spectrum is ~ 380 , dropping only by 20% to ~ 310 when corrected for the telluric absorption.

4.1.5. H_2O , O_3 , CO_2 , and CO at 4700 nm

Ozone (O_3) is a major opacity source in the UV where the Hartley and Huggins bands block most of the stellar and solar radiation. The Chappuis bands at 600 nm are weak, but still visible in optical spectra (Griffin 2005) and originate from the stratosphere where ozone has its maximum abundance at about 30 km height (see Fig. 1). The near-infrared band of O_3 , centered at 4750 nm, is produced by overtone and combination transitions. These lines are strong but comparably narrow, which is due to the low atmospheric pressure at the height of formation.

We chose a standard star observation from April 22, 2007 covering a region around 4700 nm at airmass 1.06. The resolving power was $\sim 87\,000$ at a slitwidth of $0.2''$. In the chosen region, H_2O , CO_2 , and CO lines contribute to the opacity, see Fig. 8. The Q branch of the 20001 \leftarrow 01101 transition of CO_2 is very prominent (see the second to last panel in Fig. 8).

The fundamental transitions of carbon monoxide (CO) are of great importance for astronomical observations. Unfortunately, the CO lines are also very strong in the telluric spectrum and the transitions with $2 \leq J \leq 12$ are saturated, even at

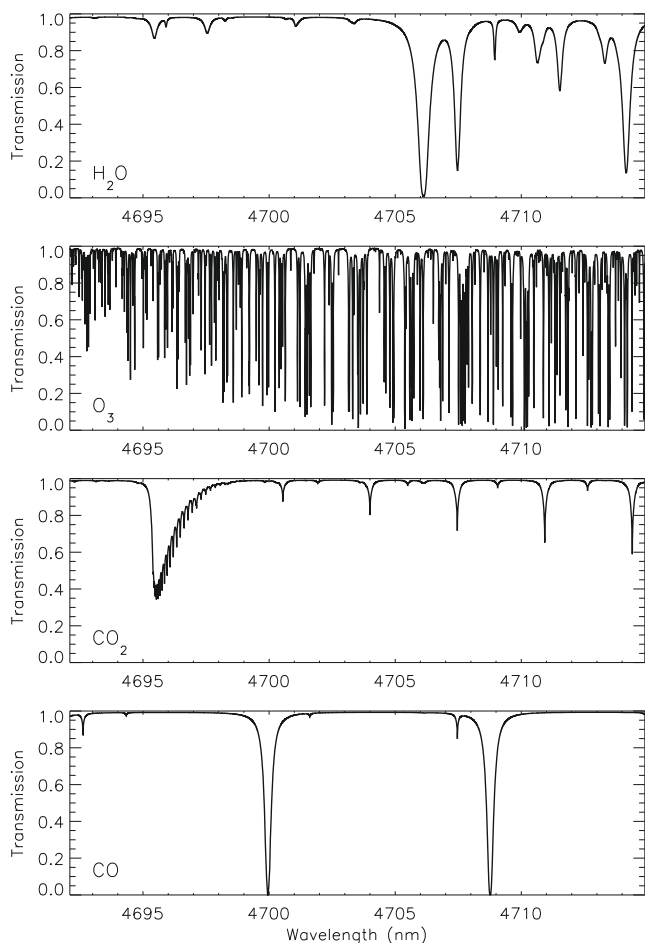


Fig. 8. Transmission of the individual species involved in the spectral region shown in Fig. 9. From top to bottom: water vapour (H_2O), ozone (O_3), carbon dioxide (CO_2), and carbene monoxide (CO). The transmissions are shown before convolution with the instrumental profile.

modest airmasses. High spectral resolution and velocity offsets from barycentric motion and radial velocity of the science target are needed to disentangle the stellar and telluric components. The average SNR of the measured spectrum in Fig. 9 is ~ 200 , which dropped to ~ 100 when divided by the model spectrum. The CO_2 Q branch is not well fitted and peaks out as a strong residual near 4695 nm. The second strongest residual near 4707.5 nm is a H_2O line. This line was much better fitted with the line data from the HITRAN 2004 database before the update in 10/2008 (AER V2.1). Its line strength is notably different in the newest line file (AER V2.2) and causes a worse fit. Other minor residuals are mostly from O_3 . Because the line formation happens in a cold part of the Earth's atmosphere at ~ 20 – 30 km height and its abundance sharply decreases below 15 km, O_3 does not contribute notably to the telluric emission received at the ground but is a strong source of absorption. Only small scaling corrections at the order of 2.5–3% had to be applied to the model abundances of CO_2 , and CO while the O_3 abundance had to be increased by 6%. The best fit PWV was 7.2 mm (7.7 mm using the HITRAN 2004 database).

4.2. Limitations

For the cases presented in the previous subsection, we have never reached the formal noise limit of the observations but we found

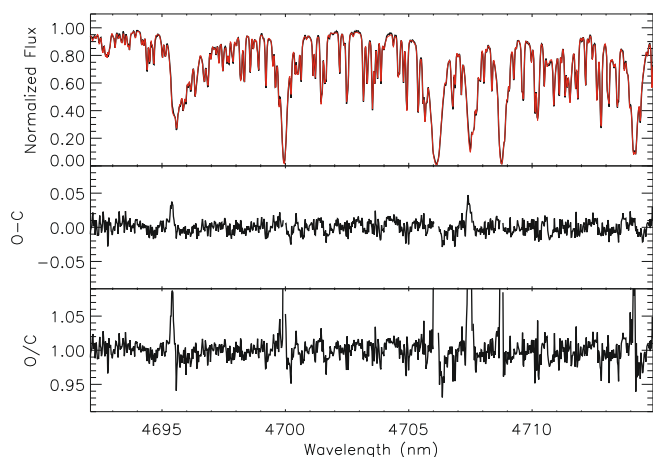


Fig. 9. Same as Fig. 4 but for ozone (O_3), water vapour (H_2O), carbon dioxide (CO_2) and carbene monoxide (CO) on chip one of CRIRES setting $\lambda_{\text{ref}} 4760.8$ nm.

a satisfying fit of the model to the data. To illustrate the limits of our approach, we present an extreme example of very badly fitting H_2O lines in Fig. 10. The airmass of the observation was 1.3, the resolving power was $\sim 64\,000$ at a slitwidth of $0.2''$. This is the worst fit we have come across and it is clearly related with the line strength information in HITRAN. Neighboring lines in the model are both too strong and too weak. The predicted PWV at zenith was ~ 2 mm and was scaled down during the optimisation to 0.35 mm. We find this reasonable, since the PWV was reported at the same level for other spectral settings taken during the same night where the water vapour lines were fitted much better. The lines in the setting shown in Fig. 10 are thus very strong lines as they show up with an average transmission of 50% and below even under such dry conditions. Water vapour lines also caused problems in the setting presented in Sect. 4.1.5, and left residuals in other settings. It is thus the water vapour which appears the hardest to model and to be affected the most from uncertainties in the line database.

4.2.1. Limitations from the HITRAN database

The HITRAN database contains estimators on the uncertainty of position, line strength and pressure broadening for all entries (Rothman et al. 2005). For example, the $^{12}\text{CO}_2$ lines shown in Fig. 5 have an average uncertainty in their positions of $\sim 5 \times 10^{-5} \text{ cm}^{-1}$ (Miller & Brown 2004) or ~ 3 m/s and the uncertainty in the line strength is at the few-% level (Tashkun et al. 2003). The strongest CH_4 lines shown in Fig. 6 have an average uncertainty in their positions of $\sim 10^{-3} \text{ cm}^{-1}$ or ~ 60 m/s and uncertainties in their line strengths of up to 10% (Rothman et al. 2005; Brown et al. 2003). This level of positional uncertainties influences the wavelength solution and the line fitting. As expected, even the strongest water vapour lines shown in Fig. 10 have positional uncertainties of up to 0.01 cm^{-1} or $\sim 2.6 \text{ km s}^{-1}$ and their line strength can be uncertain to more than 20% (Rothman et al. 2005). This explains the bad fit to the data shown in Fig. 10.

4.2.2. Limitations from LBLRTM

Other limitations arise from the limited treatment of line mixing (aka line coupling) effects and from the assumed line profile. Line mixing in LBLRTM is modeled using a first order perturbation approach. For CO_2 , line mixing is treated according to Niro et al. (2005). LBLRTM employs a Voigt line shape at all

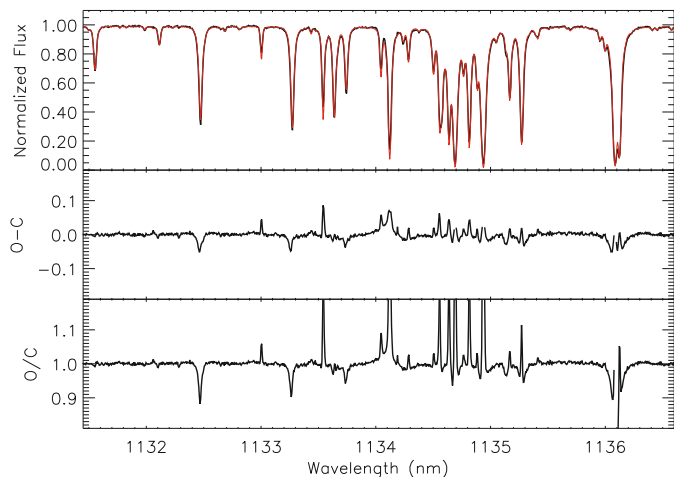


Fig. 10. Same as Fig. 4 but for water vapour (H_2O) lines on chip three of CRIRRES setting λ_{ref} 1134.0 nm.

atmospheric levels with an algorithm based on a linear combination of approximating functions (Clough et al. 2005). A Rautian, Galatry or speed-dependent Voigt profile would be needed to account more precisely for speed-dependent broadening effects and line mixing effects (e.g., Brault et al. 2003, Fig. 1). The errors induced by this simplification are, however, comparatively small against the uncertainties in other line parameters entries of the HITRAN database.

4.2.3. Limitations from the model atmosphere

As long as the total column densities of the different molecular species can be varied, the uncertainties in the model atmosphere have only a minor impact on the model spectrum. Exceptions are CO and O_3 , where changes in the vertical abundance profiles lead to notably different line widths, because either the abundance maximum or minimum of the species is close to the temperature inversion of the Earth's atmosphere (see Figs. 1 and 2). The model for our mid-infrared case presented in Sect. 4.1.5 does show a dependence on the type of model atmosphere, i.e. whether an equatorial or night-time MIPAS model is used. The intrinsic line width is, however, degenerate with the instrumental profile and while the formal resolution measured as the width of the instrumental profile varies, the residuals give no hint which of the two model atmospheres is more representative of the actual conditions.

For the water vapour, we did not find a significant improvement in the fit quality when using the MM5 atmospheric models instead of the GDAS models. The fitting delivered the same PWV for both models and the different vertical distribution of the water vapour had no noticeable influence on our test cases. Re-fitting of the vertical distribution of the water vapour will thus lead only to minor improvements, if at all, and seems only feasible when enough lines are present, preferably from different vibrational transitions.

4.2.4. Limitations from stellar lines

The fitting of a telluric model to an observed spectrum will be more complicated in the presence of blends between telluric and stellar lines. Various strategies could be used to allow for the same fitting quality as compared to the featureless early type stars used in this study. (I) The usage of a stellar model spectrum in the fitting process would significantly improve the telluric model. (II) In the thermal infrared, the telluric emission lines

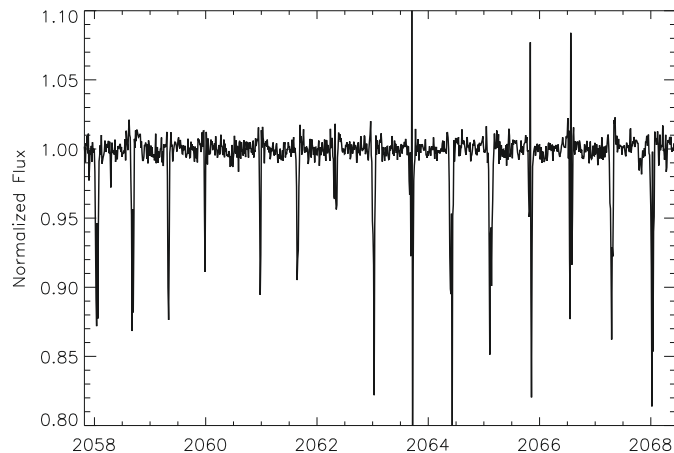


Fig. 11. Division of two spectra from consecutive nodding positions of the same star shown in Fig. 5. The spectrum from nodding position B was mapped on the wavelength scale of the spectrum from nodding position A before the division. The difference in airmass is only 0.005. See Fig. 5 and text for details.

could be used to further constrain the fit in terms of wavelength solution and molecular abundances. (III) In the worst case, an observation of a telluric standard star in the same spectral setting but not necessarily at the same time and airmass can be used to constrain the model fit for the target (e.g., Villanueva et al. 2008).

5. Conclusions

We have shown that with the use of LBLRTM, a line-by-line radiative transfer code, and an appropriate model atmosphere, synthetic telluric transmission spectra can be constructed that match the observed telluric spectrum to 2% or better (see, e.g., Fig. 5), even close to saturated lines. These results were achieved with only a modest effort in back-fitting of atmospheric properties. Compilation of LBLRTM is straight forward and when using the described IDL fitting routines for χ^2 minimisation, the time needed for a model fit to a single CRIRRES chip is less than 2 min on standard desktop PC. User interaction is only required to download the appropriate GDAS model from the NOAA website, to set up an initial wavelength scale, and to choose the species to be fitted.

Although full photon noise limited performance is never achieved, the usage of the synthetic transmission spectra offers an alternative to the usage of standard star observations. The tight constraints in matching airmass and instrumental profile of the science target imposed on telluric standard stars are hard to meet in reality without downgrading the achievable SNR or the spectral resolution.

We demonstrate this problem by using the data shown in Sect. 4.1.2, this time showing in Fig. 11 the spectrum from nodding position A divided by the one in nodding position B, taken 1.5 min later. The difference in airmass is only 0.005. The spectrum of the B position had to be mapped onto the wavelength scale of the A position due to the slit curvature which changes the wavelength zero point and the dispersion between both positions. We note that at this step the wavelength solution provided by the telluric model for both nodding positions had to be used to achieve this result. Despite the fact that both spectra are of the same star and taken close in time and airmass, a slight change in the AO performance over time and in the instrumental profile between both nodding positions led to the apparent residuals in the division. The detector response at the two nodding positions

is uniform to better than 0.5% after correcting for non-linearity and applying a flatfield correction as outlined in Sect. 2. Hence, changes in the instrumental profile are the reason for the observed residuals.

Under more realistic circumstances, i.e. when a different target is observed after a longer time interval and not matching the airmass that well, the expected residuals from a simple division by another telluric standard star would be even larger. Even in this very favorable case, the residuals in the line cores are at the order of 10% and the residuals in the wings are comparable to the ones after dividing through a telluric model (see Fig. 5). Acquiring accurate empirical telluric spectra with CRIRES is challenging and the usage of telluric model spectra is an interesting alternative to the observation of standard stars, especially since it also saves valuable observing time.

An effective removal of telluric absorption lines demands, but in most cases also provides, a wavelength calibration at an extremely high level of precision. Narrow and steep absorption lines are sensitive to wavelength errors at the level of 1/50 pixel (30 m/s). Having the telluric absorption lines directly imprinted in the science spectrum, allows for the most unbiased wavelength calibration when compared to telluric emission lines or calibration lamp lines since the telluric absorption lines experience the same IP as the stellar lines. Given the large variation in accuracy of the line positions in the HITRAN database, a wavelength solution based on a weighted fit to telluric lines offers further improvements, especially when a sufficient number of lines is present. With the excellent line positions of the $^{12}\text{CO}_2$ lines shown in Fig. 5 a wavelength solution with an rms of ≤ 20 m/s can be achieved.

To a certain extent, the Earth's atmosphere serves as a 'gas-cell' for precise wavelength calibration as shown by Seifahrt & Käufel (2008) who used N_2O lines at 4100 nm and reached a short term wavelength precision of ≤ 10 m/s over several hours.

Recently Figueira et al. (2010) have demonstrated, based on 6 years of HARPS optical data, that atmospheric O_2 absorption lines are stable down to 10 m/s and can be used as a frequency standard with an internal precision of better than 2 m/s. Transferring this precision into the infrared wavelength regime needs a precise model for the telluric interference spectrum to account for variability, mostly H_2O content and airmass. The latter is not trivial in the presence of fully or nearly fully saturated lines.

Despite not being intended as a study for time dependent water vapour distribution, we find that the observed variability in the PWV above Paranal, ranging from 0.4 mm to nearly 9 mm in our observations, exceeds the range of satellite based predictions for the PWV by far. We have compiled all predicted and fitted PWV values used in this study in Table 2. The PWV can change notably in the course of hours and re-fitting of the total PWV for a specific observation is unavoidable. A detailed study on the water vapour column density above Paranal will be presented elsewhere (Smette et al., in preparation). The GDAS meteorological model very often predicts too high PWV values, which we have always scaled globally, without changing its vertical distribution. Comparing GDAS and MM5 models for the same nights, results in the same PWV value and the use of the more realistic MM5 model did not significantly improve or fit results.

Choosing well characterised lines, we do see the potential in the CRIRES data to retrieve atmospheric abundance information for atmospheric trace gases at or below the 1% level delivering important information about the short and long term variations

of greenhouse gases from astronomical observations. Such measurements could be a valuable byproduct of astronomical spectra collected with high-resolution infrared spectrographs.

Acknowledgements. We thank all colleagues involved in the development, installation and commissioning of CRIRES. Special thanks goes to Alain Smette for his comments and suggestions on the telluric modeling and Linda Brown for valuable discussions on line broadening and -mixing. A.S. acknowledges financial support from the Deutsche Forschungsgemeinschaft under DFG RE 1664/4-1. M.J.R. and A.S. acknowledge support from the National Science Foundation under grant NSF AST-0708074. J.L.B. acknowledges research funding from the European Commissions Seventh Framework Programme as an International Incoming Fellow (grant No. PIFFGA-2009-234866).

References

- Abrams, M. C., Davis, S. P., Rao, M. L. P., Engleman, R. J., & Brault, J. W. 1994, *ApJS*, 93, 351
- Bailey, J., Simpson, A., & Crisp, D. 2007, *PASP*, 119, 228
- Bevington, P. R., & Robinson, D. K. 1992 (New York: McGraw-Hill), 2nd edn.
- Brault, J. W., Brown, L. R., Chackerian, C., et al. 2003, *J. Mol. Spectrosc.*, 222, 220
- Brown, L. R., et al. 2003, *J. Quant. Spectr. Radiat. Transf.*, 82, 219
- Clough, S. A., Kneizys, F. X., Rothman, L. S., & Gallery, W. O. 1981, *Proc. SPIE*, 277, 152
- Clough, S. A., Iacono, M. J., & Moncet, J.-L. 1992, *J. Geophys. Res.*, 97, 15761
- Clough, S. A., Shephard, M. W., Mlawer, E. J., et al. 2005, *J. Quant. Spectr. Radiat. Transf.*, 91, 233
- Dlugokencky, E. J., Houweling, S., Bruhwiler, L., et al. 2003, *Geophys. Res. Lett.*, 30, 190000
- Figueira, P., Pepe, F., Lovis, C., & Mayor, M. 2010, *A&A*, 515, A106
- Grell, G. A., Dudhia, J., & Stauffer, D. R. 1995, NCAR Tech. Note NCAR/TN-398+STR, 122
- Griffin, R. E. 2005, *PASP*, 117, 885
- Hase, F., Wallace, L., McLeod, S. D., Harrison, J. J., & Bernath, P. F. 2010, *J. Quant. Spectr. Radiat. Transf.*, 111, 521
- Horne, K. 1986, *PASP*, 98, 609
- Jacquinet-Husson, N., Arie, E., Ballard, J., et al. 1999, *J. Quant. Spectr. Radiat. Transf.*, 62, 205
- Jacquinet-Husson, N., Scott, N. A., Chedin, A., & Chursin, A. A. 2003, *Atmospher. Ocean. Opt.*, 16, 256
- Käufel, H. U., Ballester, P., Biereichel, P., et al. 2004, *SPIE*, 5492, 1218
- Käufel, H. U., Esteves, R., Fredigo, E., et al. 2006, *The Messenger*, 124, 2
- Käufel, H. U., Amico, P., Ballester, P., et al. 2006, *The Messenger*, 126, 32
- Kenworthy, M. A., & Hanson, M. M. 2004, *PASP*, 116, 97
- Kerber, F., Nave, G., & Sansonetti, C. J. 2008, *ApJS*, 178, 374
- Lallement, R., Bertin, P., Chassefiere, E., & Scott, N. 1993, *A&A*, 271, 734
- Livingston, W., & Wallace, L. 1991, NSO Technical Report, Tucson: National Solar Observatory, National Optical Astronomy Observatory
- Lord, S. D. 1992, NASA Technical Memorandum, 103957
- Miller, C. E., & Brown, L. R. 2004, *J. Mol. Spectrosc.*, 228, 329
- Niro, F., Jucks, K., & Hartmann, J.-M. 2005, *J. Quant. Spectr. Radiat. Transf.*, 95, 469
- Paufique, J., et al. 2006, [arXiv:astro-ph/0608011]
- Rothman, L. S., Jacquemart, D., Barbe, A., et al. 2005, *J. Quant. Spectr. Radiat. Transf.*, 96, 139
- Rothman, L. S., Gordon, I. E., Barbe, A., et al. 2009, *J. Quant. Spectr. Radiat. Transf.*, 110, 533
- Rousselot, P., Lidman, C., Cuby, J.-G., Moreels, G., & Monnet, G. 2000, *A&A*, 354, 1134
- Rucinski, S. 1999, in *Precise Stellar Radial Velocities*, IAU Colloq. 170, ASP Conf. Ser., 185, 82
- Scott, N. A. 1974, *J. Quant. Spectr. Radiat. Transf.*, 14, 691
- Scott, N. A., & Chedin, A. 1981, *J. Appl. Meteor.*, 20, 802
- Seifahrt, A., & Käufel, H. U. 2008, *A&A*, 491, 929
- Tashkun, S. A., Perevalov, V. I., Teffo, J.-L., Bykov, A. D., & Lavrentieva, N. N. 2003, *J. Quant. Spectr. Radiat. Transf.*, 82, 165
- Udem, T., Holzwarth, R., & Hänsch, T. W. 2002, *Nature*, 416, 233
- Vacca, W. D., Cushing, M. C., & Rayner, J. T. 2003, *PASP*, 115, 389
- Villanueva, G. L., Mumma, M. J., Novak, R. E., & Hewagama, T. 2008, *Icarus*, 195, 34
- Widemann, T., Bertaux, J.-L., Querci, M., & Querci, F. 1994, *A&A*, 282, 879
- Zängl, G. 2002, *Monthly Weather Review*, 130, 1423
- Zängl, G. 2003, *Monthly Weather Review*, 131, 2875